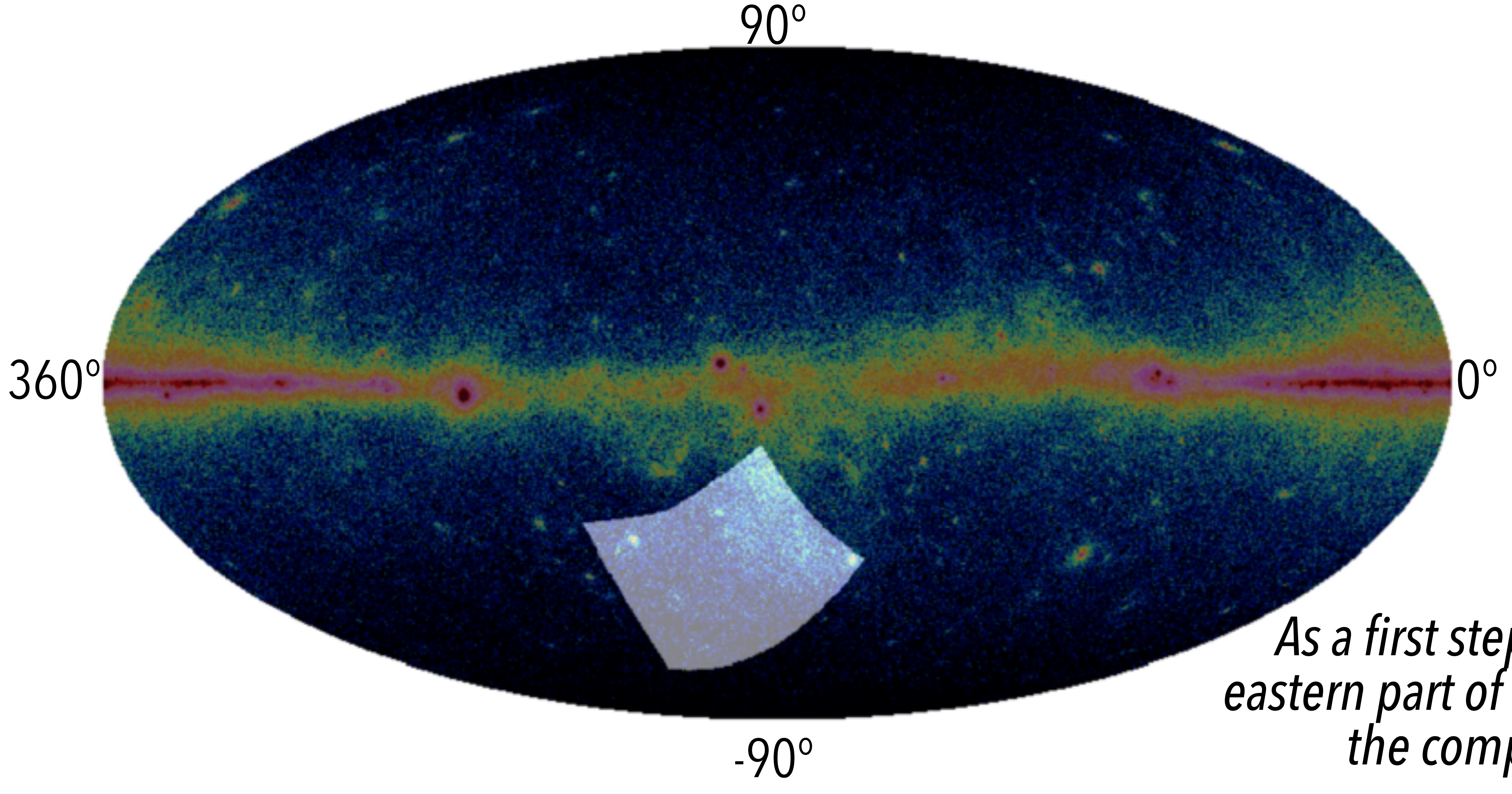
The Orion-Eridanus superbubble

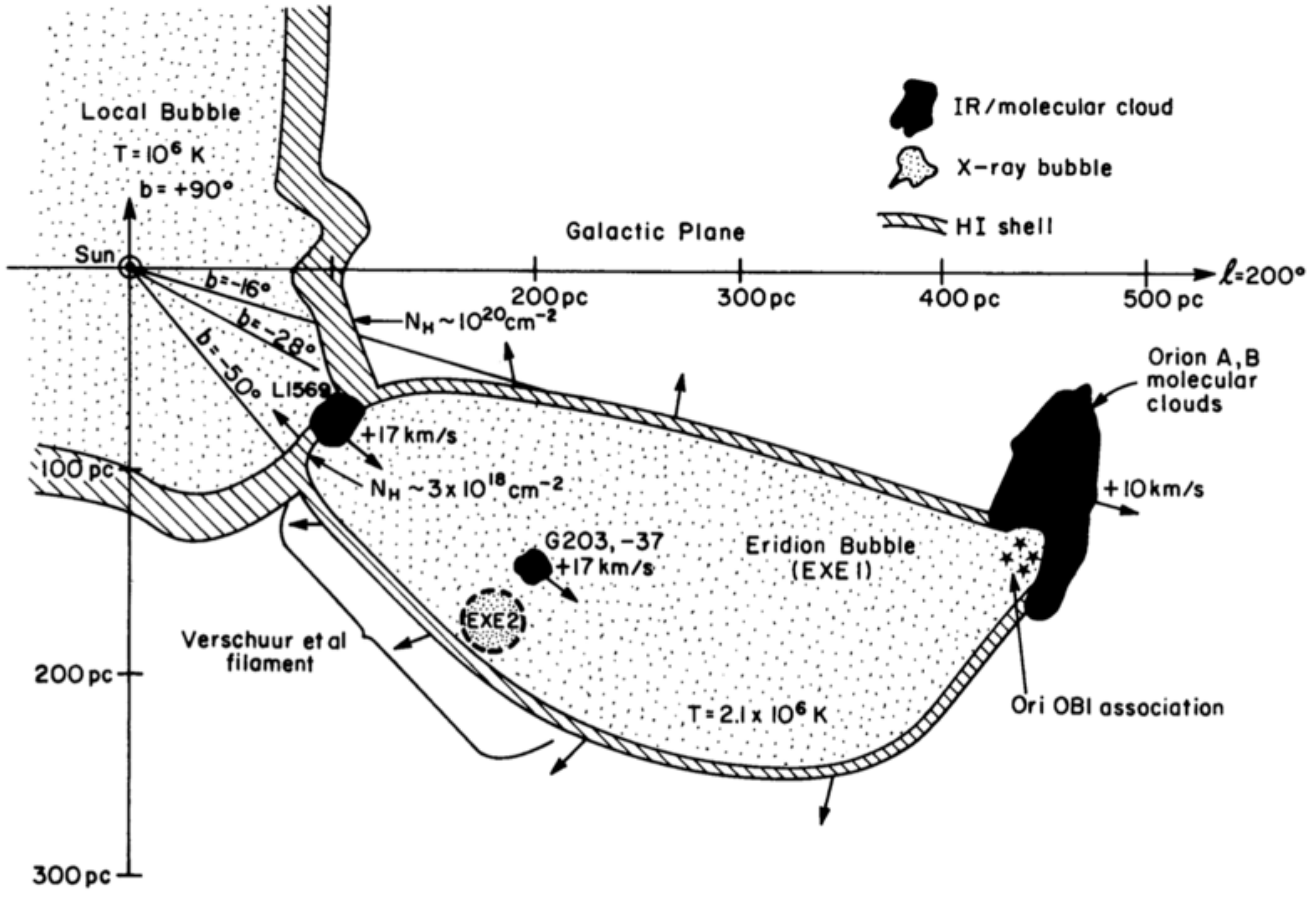
The ∼12 Myr-old superbubble [1] has been blown by the supernovae and supersonic winds of Orion's massive stars. It has likely **fostered cosmic-ray acceleration.** The large level of MHD turbulence in the bubble can also affect the diffusion properties of cosmic rays. We aim to **probe the cosmic-ray flux** inside the superbubble by comparing the γ-ray emission produced in the inner clouds with the average emissivity measured in other interstellar clouds in the solar neighbourhood [2].



As a first step, we focus on the south-eastern part of the superbubble and avoid the complex vicinity of Orion.

Characteristics of the superbubble [3]:

- last 12 Myr : 10-20 Supernovae
- last 12 Myr : 30–100 formed stars with M ≥ 8 M⊙
- d ~ 180-400 pc

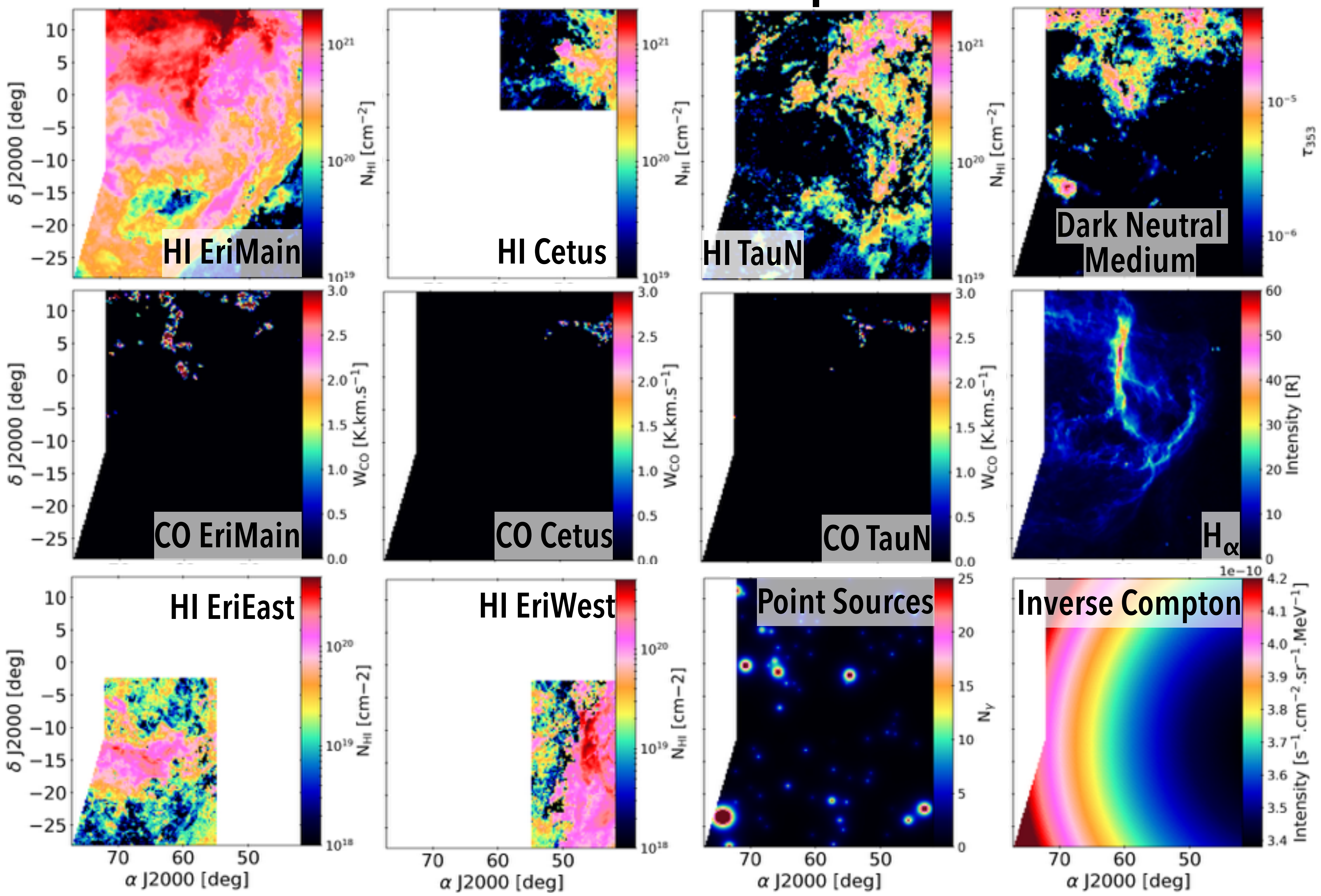


Possible geometry of the superbubble [3]

Modelling the γ-ray emission

To this aim, we have used 8 years of Fermi LAT data above 250 MeV, **the spatial and spectral distributions of which have been modelled** in terms of interstellar emission borne in the **different gas phases** of 5 clouds (atomic, dark neutral, molecular, and ionized phases). The model includes other ancillary components such as inverse-Compton emission, point sources, and solar and lunar emissions. The **atomic and molecular gas** phases are traced by radio HI and CO emission lines, the **ionized gas** by H_α optical recombination lines, and the **dark neutral medium** from the coupled analysis of the γ-rays and of the dust optical depth derived from the Planck and IRAS observations [4].

selection of model components

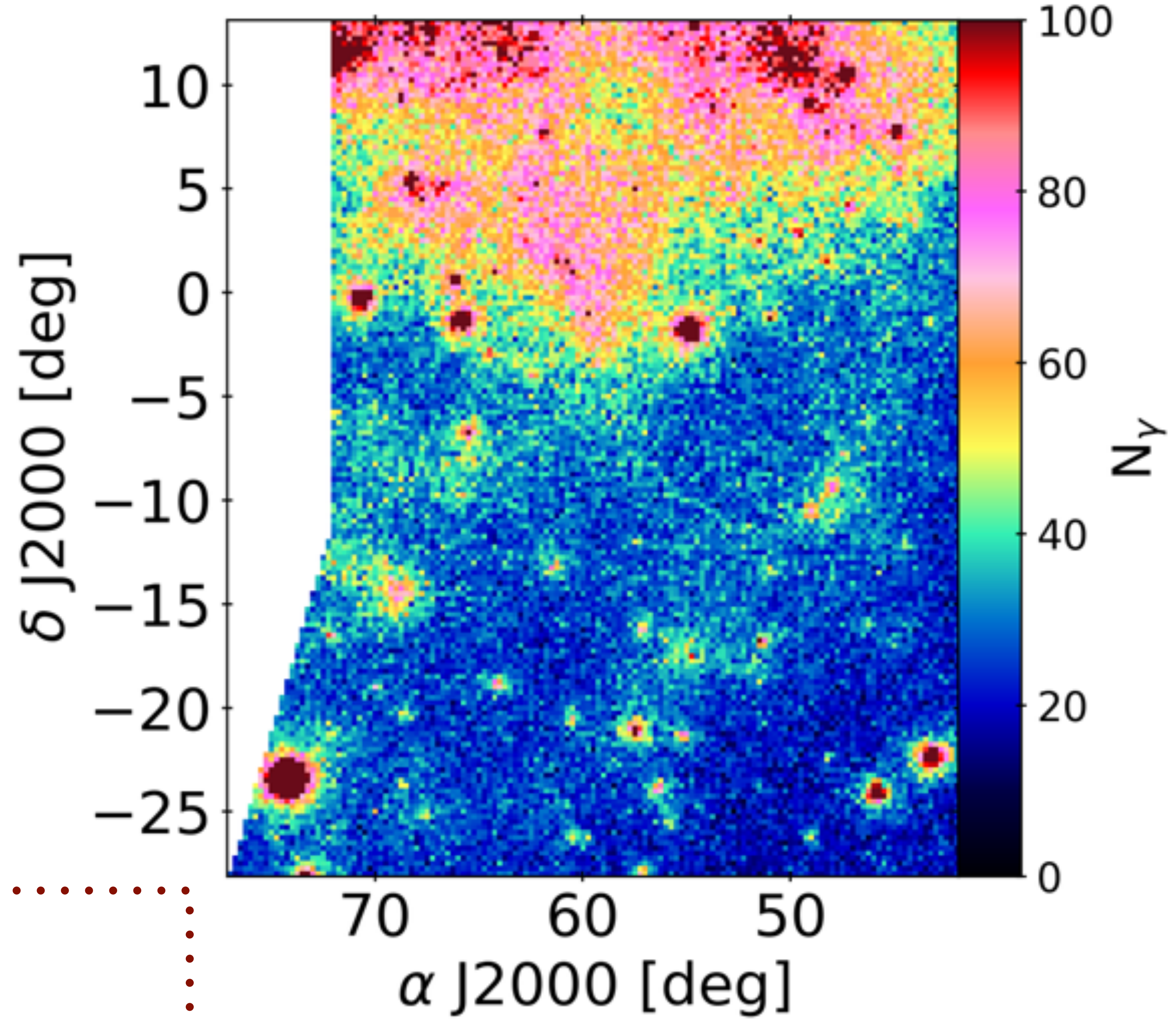


(+ Sun-Moon + an isotropic component)

FIT

- binned max. likelihood
- in 8 energy bins 0.25 GeV-TeV
- energy-dependent spatial & energy resolutions

γ rays > 250 MeV



Interstellar medium data:

HI : HI4PI (Bailin et al. 2016)

CO : CfA (Dame, 2001)

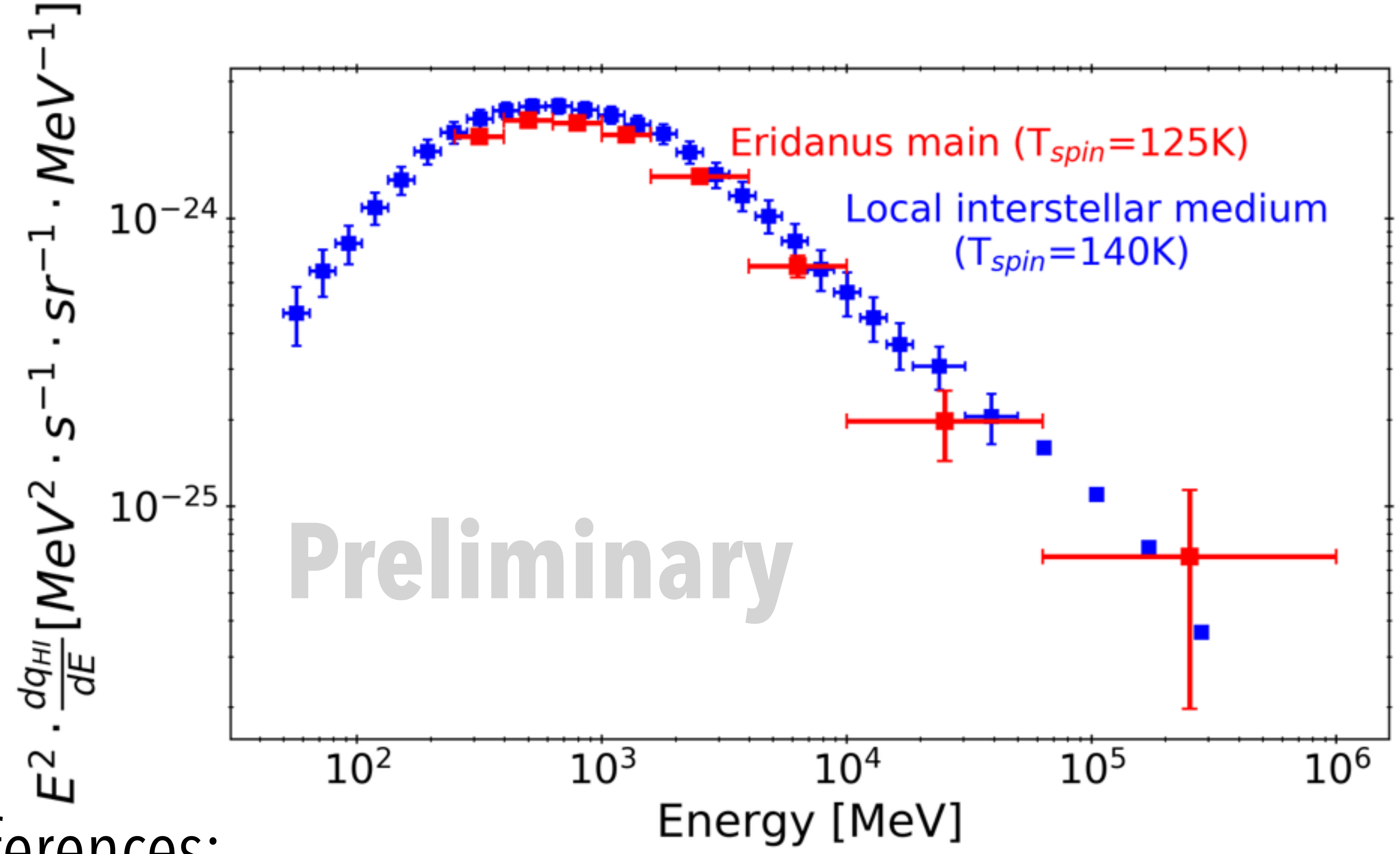
Dust : Planck + IRAS, optical depth 353GHz (Planck collab. 2016)

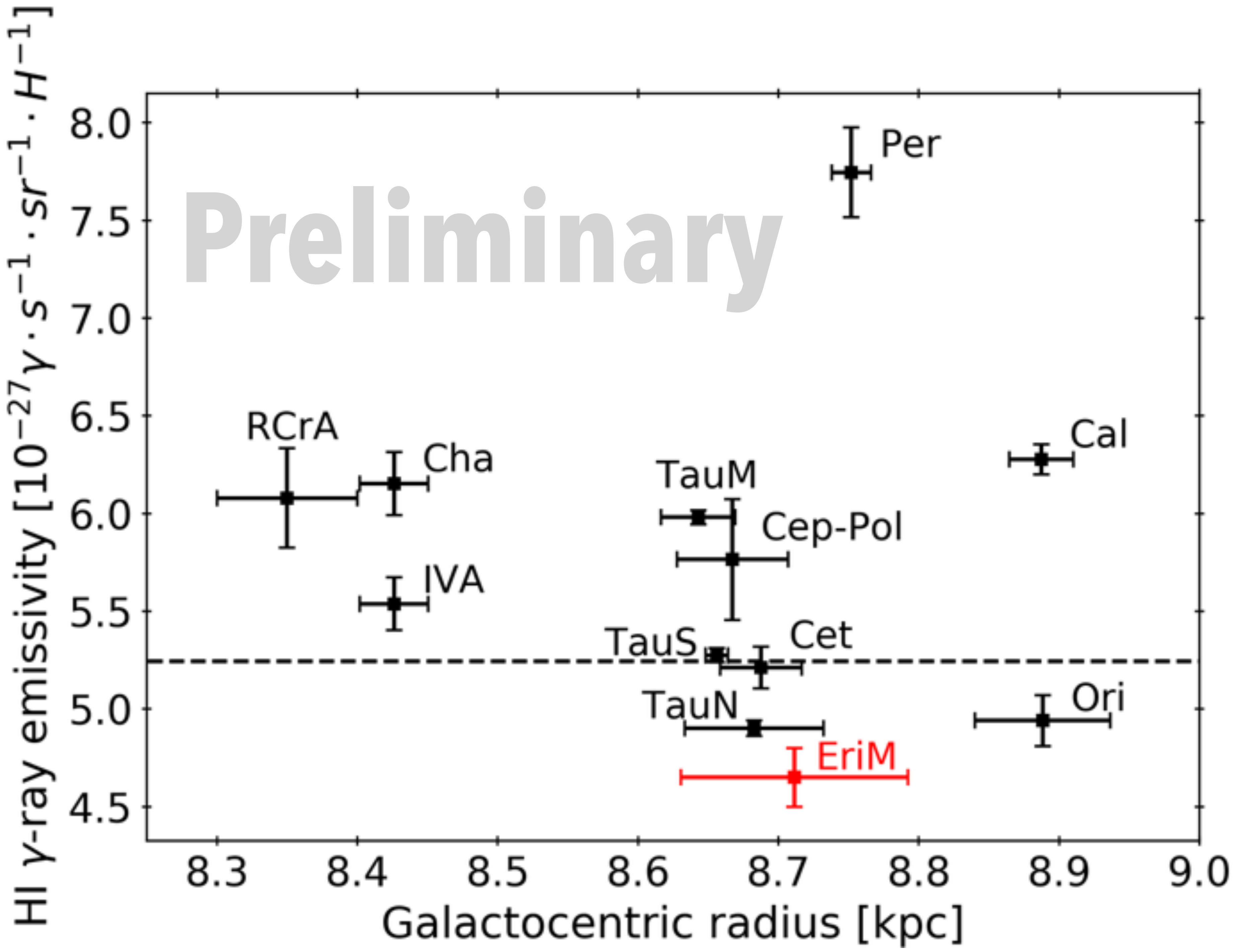
H_α : VTSS, SHASSA, WHAM (Finkbeiner 2003)

Gamma : Fermi-LAT, 8 years, 250 MeV → 1 TeV, energy-dependent cut in zmax(E) zenith angle

Results and prospects

Preliminary results show that the emissivity spectrum of the main HI cloud is **consistent with the average spectrum** measured outside the superbubble, in nearby clouds of the Gould Belt. The agreement covers the entire energy range from 250 MeV to 1 TeV, with no hint of depletion at low energies, nor of hardening at high energy. This uniformity calls for a detailed assessment of the recent **supernova rate and of the energetics of massive stellar winds** in the superbubble [5] in order to estimate the production rate and diffusion lengths of young cosmic rays and to evaluate the need, or not, to advect them away in the Gould Belt or to the halo via the local Galactic wind [6].





+: nearby clouds [7]
– –: local interstellar medium average [2]
main Eridanus cloud
all for HI spin temperatures of 125-150 K

References:

[1] J. Bally. Overview of the Orion Complex, page 459. December 2008.
[2] Jean-Marc Casandjian, *The Astrophysical Journal*, 806(2) :240, 2015.
[3] D. N. Burrows, K. P. Singh, J. A. Nousek, G. P. Garmire, and J. Good. *The Astrophysical Journal*, 406 :97–111, March 1993.
[4] Planck Collaboration, Planck intermediate results XVII, *A&A* 566 :A55, June 2014.
[5] Grenier, I. A., Gamma-ray sources as relics of recent supernovae in the nearby Gould Belt, Astronomy and Astrophysics, v.364, p.L93-L96, 12/2000
[6] R. Schlickeiser, W. R. Webber, and A. Kempf, Explanation of the Local Galactic Cosmic Ray Energy Spectra Measured by Voyager 1. I. Protons , APJ78735
[7] Q. Remy, I. A. Grenier, D.J. Marshall, and J. M. Casandjian, A&A 601, A78 (2017)